

NASA/TM—2003-212310



# Astrobiology: The Case for Venus

Geoffrey A. Landis  
Glenn Research Center, Cleveland, Ohio

---

July 2003

## The NASA STI Program Office . . . in Profile

Since its founding, NASA has been dedicated to the advancement of aeronautics and space science. The NASA Scientific and Technical Information (STI) Program Office plays a key part in helping NASA maintain this important role.

The NASA STI Program Office is operated by Langley Research Center, the Lead Center for NASA's scientific and technical information. The NASA STI Program Office provides access to the NASA STI Database, the largest collection of aeronautical and space science STI in the world. The Program Office is also NASA's institutional mechanism for disseminating the results of its research and development activities. These results are published by NASA in the NASA STI Report Series, which includes the following report types:

- **TECHNICAL PUBLICATION.** Reports of completed research or a major significant phase of research that present the results of NASA programs and include extensive data or theoretical analysis. Includes compilations of significant scientific and technical data and information deemed to be of continuing reference value. NASA's counterpart of peer-reviewed formal professional papers but has less stringent limitations on manuscript length and extent of graphic presentations.
- **TECHNICAL MEMORANDUM.** Scientific and technical findings that are preliminary or of specialized interest, e.g., quick release reports, working papers, and bibliographies that contain minimal annotation. Does not contain extensive analysis.
- **CONTRACTOR REPORT.** Scientific and technical findings by NASA-sponsored contractors and grantees.

- **CONFERENCE PUBLICATION.** Collected papers from scientific and technical conferences, symposia, seminars, or other meetings sponsored or cosponsored by NASA.
- **SPECIAL PUBLICATION.** Scientific, technical, or historical information from NASA programs, projects, and missions, often concerned with subjects having substantial public interest.
- **TECHNICAL TRANSLATION.** English-language translations of foreign scientific and technical material pertinent to NASA's mission.

Specialized services that complement the STI Program Office's diverse offerings include creating custom thesauri, building customized databases, organizing and publishing research results . . . even providing videos.

For more information about the NASA STI Program Office, see the following:

- Access the NASA STI Program Home Page at <http://www.sti.nasa.gov>
- E-mail your question via the Internet to [help@sti.nasa.gov](mailto:help@sti.nasa.gov)
- Fax your question to the NASA Access Help Desk at 301-621-0134
- Telephone the NASA Access Help Desk at 301-621-0390
- Write to:  
NASA Access Help Desk  
NASA Center for Aerospace Information  
7121 Standard Drive  
Hanover, MD 21076

NASA/TM—2003-212310



# Astrobiology: The Case for Venus

Geoffrey A. Landis  
Glenn Research Center, Cleveland, Ohio

National Aeronautics and  
Space Administration

Glenn Research Center

---

July 2003

This report is a formal draft or working paper, intended to solicit comments and ideas from a technical peer group.

Available from

NASA Center for Aerospace Information  
7121 Standard Drive  
Hanover, MD 21076

National Technical Information Service  
5285 Port Royal Road  
Springfield, VA 22100

Available electronically at <http://gltrs.grc.nasa.gov>

# **Astrobiology: The Case for Venus**

Geoffrey A. Landis  
National Aeronautics and Space Administration  
Glenn Research Center  
Cleveland, Ohio 44135  
E-mail: geoffrey.a.landis@nasa.gov

## **1. Why Astrobiology?**

The scientific discipline of astrobiology addresses one of the most fundamental unanswered questions of science: are we alone? Is there life elsewhere in the universe, or is life unique to Earth?

The field of astrobiology includes the study of the chemical precursors for life in the solar system; it also includes the search for both presently existing life and fossil signs of previously existing life elsewhere in our own solar system, as well as the search for life outside the solar system. Two of the promising environments being currently considered within the solar system are the surface of the planet Mars, and the hypothesized oceans underneath the ice covering the moon Europa. Both of these environments differ in several key ways from the environments where life is found on Earth; the Mars environment in most places too cold and at too low pressure for liquid water to be stable, and the sub-ice environment of Europa lacking an abundance of free energy in the form of sunlight.

The only place in the solar system where we know that life exists today is the Earth. To look for life elsewhere in the solar system, one promising search strategy would be to find and study the environment in the solar system with conditions that are most similar to the environmental conditions where life thrives on the Earth. Specifically, we would like to study a location in the solar system with atmospheric pressure near one bar; temperature in the range where water is liquid, 0 to 100 C; abundant solar energy; and with the primary materials required for life, carbon, oxygen, nitrogen, and hydrogen, present. Other than the surface of the Earth, the only other place where these conditions exist is the atmosphere of Venus, at an altitude of about fifty to sixty kilometers above the surface [1-4].

## **2. Venus**

### **2.1 The Suitability of the Venus Environment for Life**

In the early solar system, about four billion years before present, the sun was approximately twenty-five percent less luminous than it is today. Under these conditions, it is plausible to suggest that Venus was much more Earthlike than it is today. As the solar luminosity increased, Venus became trapped in a runaway “moist greenhouse effect.” [3] The rising temperature increased the evaporation of water vapor from the oceans; this increased water vapor in the atmosphere increased the trapping of infrared radiation, which increased the heating of the planet. Eventually this feedback loop resulted in the oceans boiling dry (and hence releasing very large amounts of greenhouse gas into the atmosphere), and shortly thereafter any surface carbonate rocks decomposed into the primary minerals plus carbon dioxide, releasing their carbon dioxide back into the atmosphere. The thick carbon dioxide atmosphere and intense

greenhouse effect results in the high atmospheric pressure (about 90 bar) and high temperature (approximately 450C) found on the surface of Venus now.

However, above the cloud layers, the atmosphere of Venus is remarkably earthlike (although lacking free oxygen). As shown in figure 1, the atmospheric pressure is one bar, equal to the pressure of Earth's surface, at an altitude of about fifty kilometers above the surface. At the altitude where the pressure is about one bar, the temperature is about 350 K; decreasing to 263 K at an altitude of 60 km above the surface.

While the surface is perpetually cloud covered (figure 2), at approximately the one bar level, considerable sunlight is available. The atmospheric composition is primarily carbon dioxide and nitrogen; the cloud aerosol particles are primarily sulfuric acid (H<sub>2</sub>SO<sub>4</sub>) droplets. The elements required for life, then, carbon, oxygen, nitrogen, and hydrogen, are all present.

## **2.2 Scientific Questions in the Study of Venus**

Venus is an interesting subject for scientific study.

The planet Venus has a number of fundamental scientific mysteries, both in its atmosphere, and on its surface. Crisp et al. [5] calls the environment of Venus "among the most enigmatic in the solar system." He notes that, although the volatile inventory of Venus is not yet fully characterized, existing measurements suggest that the relative abundances of the noble gases in its atmosphere are much more solar-like than those on the other two terrestrial planets. Similarly, although today Venus has an inventory of water that is a hundred thousand times less water than that of Earth, the deuterium to hydrogen ratio in the atmosphere is on the order of 150 larger for Venus than for Earth. Since hydrogen is more easily lost from the atmosphere than deuterium, this suggests that the original amount of water present on Venus may have been comparable or larger than that on Earth. The detailed process and timing of this large loss of water are not known, leaving a large blank area in efforts to understand the history and change of the surface and atmosphere. Likewise, the cycle of sulfur in the Venus atmosphere is not yet understood. The sulfur interacts with the surface minerals, which serve as a repository of sulfur. Are there presently active sources of atmospheric sulfur, in the form of volcanoes or gas vents?

The details of the atmospheric dynamics of Venus are also not yet understood. Although the surface of Venus rotates very slowly, with a 242 day period, the atmosphere at the cloud level and higher moves around the planet a period of 4 days, 60 times faster than the surface. The question of exactly what mechanism supports this atmospheric super-rotation is a question that has eluded detailed explanation since its discovery by spacecraft in the early 1970s.

The surface of Venus is also of great scientific interest. What has caused the geologically recent resurfacing of the planet? Venus lacks large tracts of young and old crust, a signature of Earth's plate tectonics, and lacks evident plate boundary features. The mountaintops of Venus have some indication of interesting mineralogy that is different from that of the lower altitudes. The nature of this mineralogy is unknown.

The most exciting discovery remaining to be made in the solar system would be the discovery of life outside of Earth. As a near twin of Earth, Venus is an interesting subject for astrobiological study. Astrobiological studies on Venus would focus on four questions: searches for past life; studies of present chemistry, atmospheric evolution and climate; investigating the possibility of present-day atmospheric life; and studying the possibility of future life.

### **3. Astrobiology on Venus**

#### **3.1 Past Life**

If early Venus in fact had a temperate climate and an ocean, could life have evolved on Venus? This question will be extraordinarily difficult to answer for a number of reasons. Even on the surface of the Earth, four-billion year old fossil life is difficult to find and difficult to unambiguously identify once found. The environment of Venus is more hostile than that of Earth, and the planetary exterior is resurfaced by volcanism more recently than that of Earth. However, finding and studying fossils of early Venus life, if it ever existed, would be of great interest to astrobiology.

#### **3.2 Chemistry and Evolution of Atmospheres**

Studying Venus is also significant toward understanding the chemistry of atmospheres and the evolution of terrestrial planet atmospheric chemistry, allowing us to begin to understand how planetary conditions and chemistry change from the formation of the planet to the present day. Understanding the atmosphere of Venus is important to learning about and understanding the climate and atmosphere of Venus and the Earth. It is also of significant interest to astrobiology, in order to elucidate by what mechanisms the climate of a planet will stay within the habitable “liquid water” zone, whether such planets are common or uncommon in the universe, and whether the chemical evolution of atmospheres and carbon compounds is a likely or unlikely process. If we are to search for chemical signatures of the possible building blocks of biology and the current existence of life in other solar systems, it will be necessary to first understand the variety of chemistry possible in the Earthlike planets in our own solar system, and to understand which chemical processes are unambiguous signs of life, and which can be produced by abiological means.

In addition, Venus is the best laboratory in the Solar system for study of the greenhouse effect. If we are to understand climate change on Earth, we can learn by comparing and contrasting the history and climate of the Earth with that of its nearby neighbors, Venus and Mars. A measurement of infrared absorption, reflection, and emission in the atmosphere of Venus will do much toward elucidating the greenhouse effect mechanisms on Venus, and allowing us to compare and contrast the greenhouse effect on Venus with that of Mars and Earth, and understand the effect of life and the impact of human actions on the environment

#### **3.3 Present Life**

Could bacterial life exist in the atmosphere of Venus today? Although this is considered unlikely, the possibility of life in the clouds or the middle atmosphere of Venus has not been ruled out by any observations made to date. While the atmosphere is both dry and acidic, extremophilic life has adapted to far more harsh conditions on Earth.

There is some evidence that the trace-gas constituents of the Venus atmosphere are not in chemical equilibrium with each other. On Earth, the primary source of disequilibrium in the atmospheric chemistry is the activities of biological processing; could disequilibrium on Venus also be a sign of life? In 1997, David Grinspoon made the suggestion that microbes in the clouds and middle atmosphere [4] could be the source of the disequilibrium. In 2002, Dirk Schulze-Makuch [6] independently proposed that observations of the Venus atmosphere by space probes showed signatures of possible biological activity.

As noted by Grinspoon and Schulze-Makuch, the Venus atmosphere has several trace gasses which are not in chemical equilibrium. The Venera missions and the Pioneer Venus and Magellan probes found that carbon monoxide is scarce in the planet's atmosphere, although solar radiation and lightning should produce it abundantly from carbon dioxide. Hydrogen sulfide and sulfur dioxide, two gases which react with each other and thus should not be found together, are also both present, indicating some process (possibly biological?) is producing them. Finally, although carbonyl sulfide is difficult to produce inorganically, it is present in the Venusian atmosphere. On Earth, this gas would be considered an unambiguous indicator of biological activity. While none of these chemical combinations are in themselves an unambiguous sign of life, it is interesting enough to warrant a more careful look at the atmospheric chemistry.

Another interesting sign is the nature of the ultraviolet-absorbing aerosols that form the markings seen in UV images of the planet (figure 2). The nature of these aerosols, and whether they are biological in origin, is still unknown.

On Earth, viable microorganisms are found in clouds. A research team from the University of Innsbruck examined bacteria from cloud samples by freezing onto Teflon plates water droplets collected in-situ from cloud samples from Mount Sonnblick. They then melted the samples in the laboratory and monitoring bacterial growth at the low temperatures found in clouds [7]. On the average, the cloud droplets contained around 1500 bacteria per milliliter, including round, rod-shaped and filamentary shaped bacteria, actively metabolizing at the cloud temperatures. Because the bacteria multiply over a time scale of several days, shorter than the typical lifetime of a cloud, they concluded that the bacteria collected "almost certainly reproduced inside the cloud." Cloud temperatures on Venus are similar to the range of cloud temperatures on Earth, although the Venus clouds are strongly acidic in composition. Nevertheless, identification of cloud-dwelling bacteria on Earth is a strong supporting evidence for the plausibility of bacterial life in the atmosphere of Venus.

Most of the primary materials required for life are present in easily accessible volatile form in the atmosphere of Venus. Almost 97% of the mass of living organisms (over 99.9% by atomic percentage) consists of the four elements: hydrogen, oxygen, carbon, and nitrogen; about 98% of the mass consists of these four plus sulfur and phosphorus. Of these, hydrogen, oxygen, carbon, nitrogen, and sulfur are present in significant amounts in the atmosphere. Sulfur is both a component of amino acids, and also an electron donor/acceptor anion used for respiration by some anaerobic bacteria in environments where free oxygen is not abundant; this may be a metabolic pathway that is plausible for the sulfur-rich Venus atmosphere.

The limiting resource for Venus atmospheric life would thus most likely be the requirement for phosphorus. Phosphorus is not a major constituent of the Venus atmosphere, although it is present [8]. Required cations such as calcium, potassium and sodium would also be likely to be limiting resources, along with trace mineral requirements such as magnesium and iron. It is plausible that some amount of these minerals will be present either from aeolian dust, or from volcanic sources, but these have not yet been detected in the Venus atmosphere.

Finally, the study of astrobiology has primarily been focused on life in the liquid-water phases. This is a reasonable focus, since all known life on Earth is based on carbon compounds with a liquid water medium, however, it is worth keeping in mind that other possibilities for biological chemistry may exist of which we are as yet unaware [4].



### 3.4 Future Life

As the most hospitable location in the solar system, other than Earth, the environment of Venus is worth considering as a possible future location of humans. The NASA vision, according to administrator O'Keefe, is to "Improve life here, extend life to there, and find life beyond." The possibility of terraforming Venus, first proposed by Sagan in 1961 [9], has attracted a large amount of public attention. Unfortunately, later studies [4,10,11] show that the concept of converting Venus into an Earthlike environment is a far more difficult problem than initially considered. In 1994, Sagan conceded that his concept for terraforming Venus with atmospheric microbes, then three decades old, was naïve:

"Here's the fatal flaw: In 1961, I thought the atmospheric pressure at the surface of Venus was a few bars ... We now know it to be 90 bars, so if the scheme worked, the result would be a surface buried in hundreds of meters of fine graphite, and an atmosphere made of 65 bars of almost pure molecular oxygen. Whether we would first implode under the atmospheric pressure or spontaneously burst into flames in all that oxygen is open to question. However, long before so much oxygen could build up, the graphite would spontaneously burn back into CO<sub>2</sub>, short-circuiting the process." [12].

While terraforming of Venus will not be as simple as originally suggested, it nevertheless may be possible to develop technologies to terraform the planet to habitable conditions. Possible methods of terraforming Venus are reviewed in detail by Fogg [11].

Even if Venus is not terraformed to a degree suitable for human occupation on the surface, Venus could potentially be a habitat for life bio-engineered by humans. Carl Sagan suggested that plants engineered by humans to incorporate hydrogen gas bladders could float amid the more hospitable cloud layers. In fact, hydrogen is unnecessary: even oxygen or nitrogen bladders would allow lightweight plants float in the carbon dioxide atmosphere. Kelp, for example, is a common ocean plant that uses atmosphere-filled bladders to float in the ocean (figure 3); on Venus, the same gas would be lighter than the carbon-dioxide atmosphere. While Venus may not be home to life in the present day, it may yet, in the future, have a human-engineered atmospheric ecosystem.

Finally, even if humans cannot terraform the surface of Venus, it will still be possible for humans to visit and live in aerostats in the atmosphere of Venus [13].

## 4. Venus: Spacecraft Missions

While exploration of Venus has attracted less attention than Mars exploration, a number of missions have examined our sister planet. The Russian "Venera" missions landed probes on the surface and looked down on Venus from orbit, and the "Vega" mission floated balloons into the clouds of Venus. The American Pioneer missions looked at Venus with orbiters and atmospheric probes, while Magellan imaged the topography of the surface of Venus with radar.

Japan plans to launch its first robotic mission to Venus in 2007, with the spacecraft expected to start orbiting the planet in 2009. The European Space Agency plans to launch the "Venus Express" mission, a near-duplicate of their 2003 Mars Express mission, is planned for a launch in late 2005 and arrival in 2006.

There are currently no NASA missions scheduled to revisit the planet Venus, but this might change. Several interesting Venus missions have been recently proposed, including balloon and aircraft explorations of the atmosphere of Venus [14], and sample-return missions from the surface.

## 5. Conclusion

Our sister planet, Venus, is both similar and different from the Earth, but despite differences, it is still the most earthlike of all the environments in the solar system. There is a strong case for studying astrobiology on Venus, to search for possible past life on the surface, for presently existing life in the atmosphere, to study the chemical evolution and environment of an Earthlike planet, and to speculate on the possible future of Venus as a habitat for humans.

Regardless of whether life is, or has been, present on the surface or in the atmosphere of Venus, we will learn about the atmosphere, climate, and ecosystem of the Earth by studying Venus.

## References

1. S. Vougher, D. Hunten, and R. Phillips, eds., *Venus II*, University of Arizona Press, 1997.
2. D.M. Hunten, L. Colin, T.M. Donahue, and V.I. Moroz, eds., *Venus*, University of Arizona Press, 1983.
3. J.B. Pollack, "Atmospheres of the Terrestrial Planets," in *The New Solar System*, Third Edition, Beatty and Chaikin (eds.), Cambridge University Press, 1990, pp. 91–103.
4. D.H. Grinspoon, "Chapter 6: Life on Venus: A Barren World?" *Venus Revealed: A New Look Below the Clouds of Our Mysterious Twin Planet*, Perseus Publishing (1997).
5. D. Crisp *et al.*, "Divergent Evolution Among Earth-like Planets: The Case for Venus Exploration," Planetary Decadal Study Community White Paper, Solar System Exploration Survey, 2003–2013 (2001).
6. D. Schulze-Makuch, 2nd European Workshop on Astrobiology, Graz, Austria, 16–19 Sept 2002. Reported in *New Scientist*, 25th September 2002.
7. B. Sattler, H. Puxbaum, and R. Psenner, "Bacterial growth in supercooled cloud droplets," *Geophys. Res. Lett.* Vol. 28, No. 2, pp. 239–242 (2001).
8. L.W. Esposito, J-L. Bertaux, V. Krasnopolsky, V.I. Moroz, and L.V. Zasova, "Chemistry of the Lower Atmosphere of Venus," 430–431, in *Venus II*, University of Arizona Press, S. Vougher, D. Hunten, and R. Phillips, eds., 1997.
9. C. Sagan, "The Planet Venus," *Science*, 133, 849–858, (1961).
10. S.L. Gillett, "Second Planet, Second Earth," *Analog*, 104 (12), 64–78, Dec. 1984.
11. M.J. Fogg, *Terraforming: Engineering Planetary Environments*, SAE Press 1995.
12. C. Sagan, *Pale Blue Dot: A Vision of the Human Future in Space*, Random House; November 1994.
13. G.A. Landis, "Colonization of Venus," Conference on Human Space Exploration, Space Technology & Applications International Forum, Albuquerque NM, USA, Feb. 2–6 2003. AIP Conference Proc., 654, pp. 1193–1198.
14. G.A. Landis, C. Lamarre and A. Colozza, "Atmospheric Flight on Venus," *NASA Technical Memorandum 2002-211467*; paper AIAA-2002-0819, AIAA 40th Aerospace Sciences Meeting, Reno NV, USA, January 14–17, 2002.

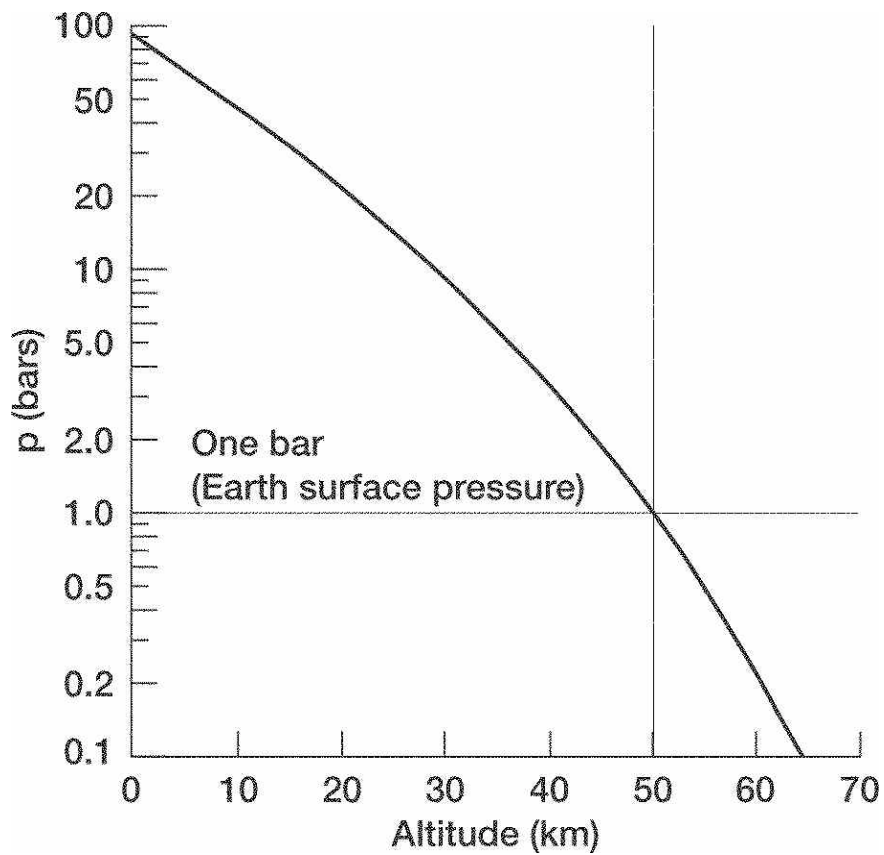


Figure 1: Atmospheric pressure as a function of altitude on the planet Venus, with the atmospheric pressure at the surface of the Earth, one bar, indicated.

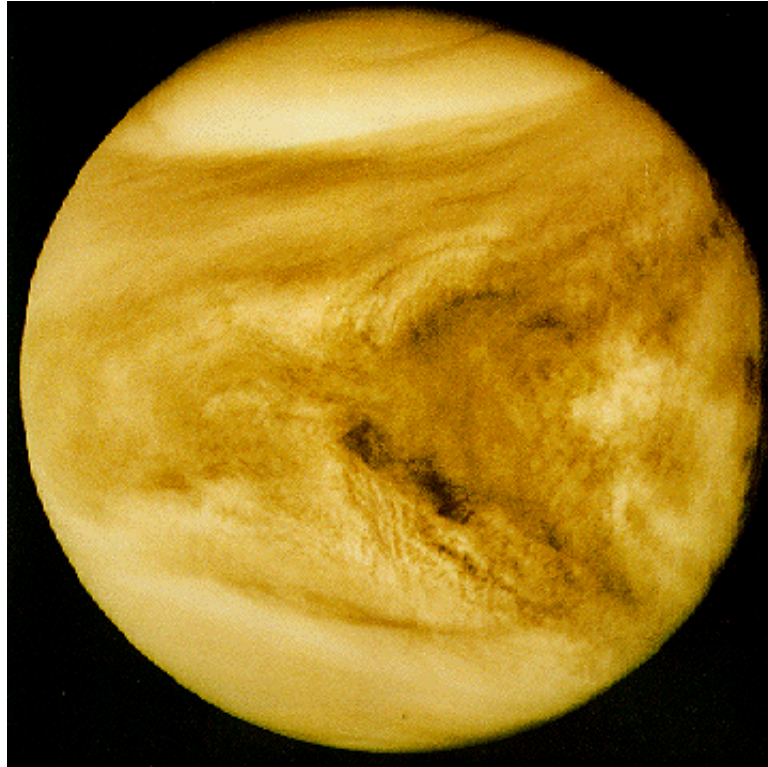


Figure 2: Venus viewed in the ultraviolet from the Pioneer spacecraft. The identity of the dark ultraviolet-absorbing aerosols forming the markings visible in the atmosphere is still unknown.



Figure 3: Kelp air bladder. The bladder gives kelp the buoyancy to float in the ocean. On Venus, the gas filling these bladders would be less dense than the carbon dioxide atmosphere (photo courtesy of Paul Kemp; used with permission).

<b>REPORT DOCUMENTATION PAGE</b>			<i>Form Approved</i> <i>OMB No. 0704-0188</i>	
Public reporting burden for this collection of information is estimated to average 1 hour per response, including the time for reviewing instructions, searching existing data sources, gathering and maintaining the data needed, and completing and reviewing the collection of information. Send comments regarding this burden estimate or any other aspect of this collection of information, including suggestions for reducing this burden, to Washington Headquarters Services, Directorate for Information Operations and Reports, 1215 Jefferson Davis Highway, Suite 1204, Arlington, VA 22202-4302, and to the Office of Management and Budget, Paperwork Reduction Project (0704-0188), Washington, DC 20503.				
<b>1. AGENCY USE ONLY (Leave blank)</b>		<b>2. REPORT DATE</b> July 2003	<b>3. REPORT TYPE AND DATES COVERED</b> Technical Memorandum	
<b>4. TITLE AND SUBTITLE</b>  Astrobiology: The Case for Venus			<b>5. FUNDING NUMBERS</b>  WBS-22-755-60-02	
<b>6. AUTHOR(S)</b>  Geoffrey A. Landis				
<b>7. PERFORMING ORGANIZATION NAME(S) AND ADDRESS(ES)</b>  National Aeronautics and Space Administration John H. Glenn Research Center at Lewis Field Cleveland, Ohio 44135-3191			<b>8. PERFORMING ORGANIZATION REPORT NUMBER</b>  E-13895	
<b>9. SPONSORING/MONITORING AGENCY NAME(S) AND ADDRESS(ES)</b>  National Aeronautics and Space Administration Washington, DC 20546-0001			<b>10. SPONSORING/MONITORING AGENCY REPORT NUMBER</b>  NASA TM-2003-212310	
<b>11. SUPPLEMENTARY NOTES</b>  Prepared for the Journal of The British Interplanetary Society, vol. 56, no. 7/8, July-August 2003, pp. 250-254. Responsible person, Geoffrey A. Landis, organization code 5410, 216-433-2238.				
<b>12a. DISTRIBUTION/AVAILABILITY STATEMENT</b>  Unclassified - Unlimited Subject Category: 91  Available electronically at <a href="http://gltrs.grc.nasa.gov">http://gltrs.grc.nasa.gov</a> This publication is available from the NASA Center for AeroSpace Information, 301-621-0390.			<b>12b. DISTRIBUTION CODE</b>	
<b>13. ABSTRACT (Maximum 200 words)</b>  The scientific discipline of astrobiology addresses one of the most fundamental unanswered questions of science: are we alone? Is there life elsewhere in the universe, or is life unique to Earth? The field of astrobiology includes the study of the chemical precursors for life in the solar system; it also includes the search for both presently existing life and fossil signs of previously existing life elsewhere in our own solar system, as well as the search for life outside the solar system. Two of the promising environments within the solar system being currently considered are the surface of the planet Mars, and the hypothesized oceans underneath the ice covering the moon Europa. Both of these environments differ in several key ways from the environments where life is found on Earth; the Mars environment in most places too cold and at too low pressure for liquid water to be stable, and the sub-ice environment of Europa lacking an abundance of free energy in the form of sunlight. The only place in the solar system where we know that life exists today is the Earth. To look for life elsewhere in the solar system, one promising search strategy would be to find and study the environment in the solar system with conditions that are most similar to the environmental conditions where life thrives on the Earth. Specifically, we would like to study a location in the solar system with atmospheric pressure near one bar; temperature in the range where water is liquid, 0 to 100 °C; abundant solar energy; and with the primary materials required for life, carbon, oxygen, nitrogen, and hydrogen, present. Other than the surface of the Earth, the only other place where these conditions exist is the atmosphere of Venus, at an altitude of about fifty kilometers above the surface.				
<b>14. SUBJECT TERMS</b>  Venus			<b>15. NUMBER OF PAGES</b> 15	
			<b>16. PRICE CODE</b>	
<b>17. SECURITY CLASSIFICATION OF REPORT</b> Unclassified	<b>18. SECURITY CLASSIFICATION OF THIS PAGE</b> Unclassified	<b>19. SECURITY CLASSIFICATION OF ABSTRACT</b> Unclassified	<b>20. LIMITATION OF ABSTRACT</b>	