

Plasma Diagnostics of Active Region Evolution and Implications for Coronal Heating

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ABSTRACT

A detailed study is presented of the decaying solar active region NOAA 10103 observed with the Coronal Diagnostic Spectrometer (CDS), the Michelson Doppler Imager (MDI) and the Extreme-ultraviolet Imaging Telescope (EIT) on-board the *Solar and Heliospheric Observatory (SOHO)*. Electron density maps formed using Si X(356.03 Å/347.41 Å) show that the density varies from $\sim 10^{10} \text{ cm}^{-3}$ in the active region core, to $\sim 7 \times 10^8 \text{ cm}^{-3}$ at the region boundaries. Over the five days of observations, the average electron density fell by $\sim 30\%$. Temperature maps formed using Fe XVI(335.41 Å)/Fe XIV(334.18 Å) show electron temperatures of $\sim 2.34 \times 10^6 \text{ K}$ in the active region core, and $\sim 2.10 \times 10^6 \text{ K}$ at the region boundaries. Similarly to the electron density, there was a small decrease in the average electron temperature over the five day period. The radiative, conductive, and mass flow losses were calculated and used to determine the

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resultant heating rate (P_H). Radiative losses were found to dominate the active region cooling process. As the region decayed, the heating rate decreased by almost a factor of five between the first and last day of observations. The heating rate was then compared to the total unsigned magnetic flux ($\Phi_{tot} = \int dA|B_z|$), yielding a power-law of the form $P_H \sim \Phi_{tot}^{0.81 \pm 0.32}$. This result suggests that waves rather than nanoflares may be the dominant heating mechanism in this active region.

Subject headings: Sun: activity – Sun: corona – Sun: evolution – Sun: UV radiation

1. INTRODUCTION

Since the discovery of highly ionized species of iron in the solar corona in the 1930’s, physicists have been puzzled by the high temperatures observed in the outer solar atmosphere (Edlen 1937). It is widely accepted that the magnetic field plays a fundamental role in the heating process, but precise measurements of the coronal magnetic field are currently impossible. Indirect methods are therefore adopted which rely on measurable quantities such as the electron temperature, density and photospheric magnetic flux. These measurements can then be compared to theoretical predictions.

Models for coronal heating typically belong to one of two broad categories. In wave (AC) heating, the large-scale magnetic field essentially acts as a conduit for small-scale, high-frequency Alfvén waves propagating into the corona. For constant Alfvén wave amplitude $\langle v^2 \rangle$, the total power dissipated in an active region is,

$$P_H = \sqrt{\frac{\rho}{4\pi}} \langle v^2 \rangle \Phi_{tot} \quad \text{ergs s}^{-1}, \quad (1)$$

where ρ is the mass density, and Φ_{tot} is the total unsigned magnetic flux,

$$\Phi_{tot} = \int dA |B_z| \quad \text{Mx}, \quad (2)$$

where B_z is the longitudinal component of the magnetic field.

In stress (DC) heating, the coronal magnetic field stores energy in the form of electric currents until it can be dissipated, e.g., by nanoflares (Parker 1988). The total power can be estimated by,

$$P_H \sim |v| \Phi_{tot}^2 \quad \text{ergs s}^{-1}, \quad (3)$$

and the constant of proportionality describes the efficiency of magnetic dissipation, which might involve the random footpoint velocity, v (Parker 1983), or simply the geometry (Browning, Sakuria, & Priest 1986; Fisher et al. 1998).

Several authors have linked the photospheric magnetic flux to EUV and X-ray line intensity. Gurman et al. (1974) found that the line intensity of Mg X (624.94 Å) was proportional to the magnetic flux density. Schrijver (1987) related the integrated intensities of chromospheric, transition region, and coronal lines to the total magnetic flux by a power law, the index of which was dependent on the scale height. This result was later confirmed by Fisher et al. (1998), who showed that X-ray luminosity is highly correlated with the total unsigned magnetic flux. Van Driel-Gesztelyi et al. (2003) also showed a power-law relationship between the mean X-ray flux, temperature, and emission measure, and the mean magnetic field by studying the long-term evolution of an active region over several rotations, at times when there were no significant brightenings.

In this paper, we study the evolution of a decaying active region using the diagnostic capabilities of the Coronal Diagnostic Spectrometer (CDS; Harrison et al. 1995) and Michelson Doppler Imager (MDI; Scherrer et al. 1995) onboard *SOHO*. Using the temperatures, densities, and dimensions of the active region, the heating rate is calculated and compared to the

total unsigned magnetic flux. These results can be put in the context of theoretical models. Section 2 gives a brief overview of the active region, a summary of the instruments involved, and a description of the data analysis techniques. Our results are given in Section 3, and discussion and conclusions in Section 4.

2. OBSERVATIONS AND DATA ANALYSIS

NOAA 10103¹ was observed by CDS, EIT, and MDI for five consecutive days during 2002 September 10–14. Figure 1 shows the general evolution of the region over that period.

2.1. The Coronal Diagnostic Spectrometer

EUV spectra were obtained with the CDS instrument, which is a dual spectrometer that can be used to obtain images with a spatial resolution of $\sim 8''$. The Normal Incidence Spectrometer (NIS), used in this study, is a stigmatic spectrometer which forms images by moving the solar image across the slit using a scan mirror. The spectral ranges of NIS (308–381 Å and 513–633 Å) include emission lines formed over a wide range of temperatures, from $\sim 10^4$ K at the upper chromosphere, through the transition region, to $\sim 10^6$ K at the corona. The details of the *AZ_DDEP1* observing sequence used in this study can be found in Table 1.

The raw CDS data were cleaned to remove cosmic rays, and calibrated to remove the CCD readout bias and convert the data into physical units of photons $\text{cm}^{-2} \text{s}^{-1} \text{arcsec}^{-2}$. Due to the broadened nature of post-recovery CDS spectra, the emission lines were fitted with modified Gaussian profiles as described by Thompson (1999). The Gaussian term was

¹See <http://www.solarmonitor.org/20020910/0103.html>

Table 1: DETAILS OF THE CDS *AZ_DDEP1* OBSERVING SEQUENCE

Parameter	Value
Date	2002 September 10–14
Region name	NOAA 10103
Instrument	CDS/NIS1
Wavelength range (Å)	332–368
Slit size (arcsec ²)	4.064×240
Area imaged (arcsec ²)	243.84×240
Exposure time (s)	50
Number of slit positions	60

defined as,

$$G(\lambda) = \exp \left[-\frac{1}{2} \left(\frac{\lambda - \lambda_0}{\sigma} \right)^2 \right], \quad (4)$$

and the wings by,

$$W(\lambda) = \frac{1}{\left(\frac{\lambda - \lambda_0}{\sigma'} \right)^2 + 1}, \quad (5)$$

where λ is the wavelength, λ_0 is the central wavelength of the line, σ is the Gaussian width, and σ' is the FWHM.

The combined function describing the line profile can be expressed as,

$$B(\lambda) = I[(1 - a)G(\lambda) + aW(\lambda)], \quad (6)$$

where I is the amplitude of the line profile and a can be the amplitude of the red or the blue wing. The line flux is then given by,

$$\begin{aligned} \int B(\lambda) d\lambda &= I\sigma \left[\left(1 - \frac{a_{red}}{2} \left(1 + \frac{a_{blue}}{a_{red}} \right) \right) \sqrt{2\pi} \right. \\ &\quad \left. + a_{red} \left(1 + \frac{a_{blue}}{a_{red}} \right) \pi \sqrt{2 \ln(2)} \right]. \end{aligned} \quad (7)$$

Table 2: IONS, FORMATION TEMPERATURES, TRANSITIONS AND WAVELENGTHS OF EMISSION LINES IDENTIFIED IN THIS WORK

Ion	Log T_e	Transition	$\lambda/\text{\AA}$
Si X	6.1	$2s^2 2p^2 \ ^3P_{1/2} - 2s 2p^2 \ ^2D_{3/2}$	347.41
Si X	6.1	$2s^2 2p \ ^3P_{3/2} - 2s 2p^2 \ ^2D_{3/2,5/2}$	356.03
Fe XIV	6.3	$3s^2 3p \ ^2P_{1/2} - 3s 3p^2 \ ^2D_{3/2}$	334.18
Fe XVI	6.4	$3s \ ^2S_{1/2} - 3p \ ^2P_{3/2}$	335.41
Fe XVII	6.7	$2p^5 \ 3s \ ^3P_1 - 2p^5 \ 3p \ ^1D_2$	347.85

These broadened Gaussian profiles were then fitted to emission lines in three wavelength intervals using the *xcfit* routine in the CDS branch of the SolarSoftWare tree (SSW; Freeland & Handy 1998; see Figure 2). Two of the intervals were centered on each of the density sensitive Si X (347.41 \AA) and Si X (356.03 \AA) lines, while the third contained the temperature sensitive Fe XIV (334.18 \AA) and Fe XVI (335.41 \AA) pair. Due to the relatively low intensity of Fe XIV (334.18 \AA) compared to that of the adjacent Fe XVI (335.41 \AA) transition, an upper constraint on the width of $\sigma = 0.4 \text{ \AA}$ (FWHM = 0.17 \AA) was placed on the Fe XIV line. The primary lines, their formation temperatures, transitions, and rest wavelengths are given in Table 2.

2.2. The Michelson Doppler Interferometer

Magnetic field measurements were taken by the MDI instrument, which images the Sun on a 1024×1024 pixel CCD camera through a series of increasingly narrow filters. The final elements, a pair of tunable Michelson interferometers, enable MDI to record filtergrams with a FWHM bandwidth of 94 m \AA . Several times each day, polarizers are inserted to measure

the line-of-sight magnetic field. In this paper, 5-minute-averaged magnetograms of the full disk were used, with a 96 minute cadence and a pixel size of $2''$.

Berger & Lites (2003) analyzed Advanced Stokes Polarimeter (ASP) and MDI magnetograms, and found that MDI underestimates the magnetic flux densities by a factor of 1.45 for values below ~ 1200 G. For flux densities higher than ~ 1200 G, this underestimation becomes nonlinear, with the MDI fluxes saturating at ~ 1300 G. Values below ~ 1200 G were therefore corrected by multiplying by 1.45, and values above ~ 1200 G were approximately corrected by multiplying by a factor of 1.9 (Green et al. 2003).

Before calculating the total unsigned magnetic flux, two final corrections were applied to the data. The first results from the fact that the measured line-of-sight flux deviates more and more from a radial measurement as one approaches the limb. For simplicity, we assume that magnetic fields in the photosphere are predominantly radial. Then, the radial field strength becomes equal to the line-of-sight field strength times $1/\cos\theta$, where θ is the heliocentric distance of the region from Sun center.

Active region areas, A , were calculated by counting all pixels above 500 G and multiplying by the appropriate factor to obtain the active region area in cm^2 . This threshold value for the magnetic field was found to adequately separate active region structures from neighbouring areas of quiet-Sun and plage. As the region approached the limb, the effects of foreshortening became significant. Measured areas were therefore corrected by dividing by $\cos\theta$. The resulting total unsigned flux was then calculated using Equation (2). Active region volumes and lengths used later in the paper were calculated assuming that the region occupied a hemispherical volume (i.e., $V = 2/3 \pi^{-1/2} A^{3/2}$ and $L = \sqrt{\pi A}$).

3. RESULTS

3.1. Morphology and Magnetic Field

The top row of Figure 1 shows a series of $360'' \times 360''$ EIT images obtained in the 195 Å bandpass. The first three images show a number of loops to the south and north of the region, which are not visible from 2002 September 13 onwards. On September 10, the MDI magnetogram shows a simple bipolar β region, which is then classified as a $\beta\gamma$ on the following day. The region subsequently decreased in both size and complexity as it approached the west limb. The overall decay of the active region is clearest in the MDI continuum images in the bottom row of Figure 1.

Table 3 and Figure 3 show the decay of the region in terms of the cosine-corrected area and magnetic flux for September 10–14. The region was observed to have an initial area of 1.26×10^{19} cm², which fell to 6.54×10^{18} cm² by September 14. The total unsigned magnetic flux density also shows a similar trend, falling from close to 4.80×10^5 Mx cm² to 1.40×10^5 Mx cm². The product of the region area times the total magnetic flux density is then given in the bottom panel of Figure 3. As the region decays, the total unsigned magnetic flux falls off by a factor of 5–6 over the five days from September 10 to 14.

3.2. Electron Densities

Electron density maps were generated using the Si X (356.03 Å/347.41 Å) ratio in conjunction with theoretical data from the CHIANTI v4.2 atomic database (Dere et al. 1997). Figure 4 shows a plot of the density sensitive Si X (356.03 Å/347.40 Å) ratio together with the expected range of densities in an active region of this size and class.

The density maps, presented in the second row of Figure 5, show values of $\sim 10^{10}$ cm⁻³

Table 3: PRINCIPAL PARAMETERS DERIVED FOR NOAA 10103 FOR 2002 SEPTEMBER 10–14

	$\langle N_e \rangle$	$\langle T_e \rangle$	$\langle P_e \rangle$	A	V	L
Date	10^9cm^{-3}	10^6K	10^{15}K cm^{-3}	10^{18}cm^2	10^{28}cm^3	10^9cm
10 September 2002	2.46 ± 0.31	2.26 ± 0.03	3.10 ± 0.60	12.66 ± 0.75	1.69 ± 0.15	6.31 ± 0.37
11 September 2002	2.71 ± 0.53	2.24 ± 0.03	3.41 ± 0.84	10.23 ± 0.64	1.23 ± 0.11	5.66 ± 0.35
12 September 2002	3.04 ± 0.79	2.26 ± 0.03	3.83 ± 1.16	10.58 ± 0.76	1.29 ± 0.14	5.76 ± 0.41
13 September 2002	2.46 ± 0.52	2.24 ± 0.02	3.10 ± 0.65	9.89 ± 1.03	1.17 ± 0.18	5.57 ± 0.58
14 September 2002	1.87 ± 0.20	2.16 ± 0.02	2.35 ± 0.43	6.54 ± 1.22	0.62 ± 0.17	4.53 ± 0.84

	P_R	P_C	P_F	P_H	Φ_{tot}
Date	$10^{25} \text{ergs s}^{-1}$	$10^{25} \text{erg s}^{-1}$	$10^{25} \text{erg s}^{-1}$	$10^{25} \text{ergs s}^{-1}$	10^{24}Mx
10 September 2002	3.37 ± 0.90	0.28 ± 0.07	0.25 ± 0.04	3.90 ± 1.57	5.97 ± 0.35
11 September 2002	2.97 ± 1.12	0.19 ± 0.05	0.37 ± 0.08	3.53 ± 1.92	4.01 ± 0.25
12 September 2002	3.92 ± 2.07	0.21 ± 0.06	0.43 ± 0.12	4.56 ± 3.10	3.97 ± 0.28
13 September 2002	2.33 ± 1.05	0.30 ± 0.07	0.42 ± 0.14	2.59 ± 1.40	2.78 ± 0.29
14 September 2002	0.71 ± 0.24	0.07 ± 0.03	0.18 ± 0.06	0.97 ± 0.68	0.91 ± 0.17

in the active region core, and $\sim 7 \times 10^8 \text{ cm}^{-3}$ in the region boundaries. These are in good agreement with previous active region density measurements (Gallagher et al. 2001; Warren & Winebarger 2003). Densities in the core were observed to fall from $\sim 10^{10}$ on September 10 to $3.9 \times 10^9 \text{ cm}^{-3}$ on September 14, excluding the high density limit ($\sim 2.5 \times 10^{10} \text{ cm}^{-3}$) that was reached during a C-class flare on September 13. This behaviour is more evident in the top panel of Figure 6, which clearly shows that the average electron density changed by $\sim 30\%$ over the five days. Subsequently, plasma that was found to reach the high density limit was excluded from any further calculations. The average electron densities are listed in Table 3.

3.3. Electron Temperatures

Electron temperature maps were created using the method of Brosius et al. (1996), which employs the ratio of the fluxes of various ionization stages of iron. This method obtains a polynomial fit to the logarithm of the temperature as a function of the logarithm of the emissivity ratio for selected line pairs, under the assumption of an isothermal plasma, with the form,

$$\log T_e = a_0 + a_1(\log R) + a_2(\log R)^2 + a_3(\log R)^3, \quad (8)$$

where the parameters, a_0 , a_1 , a_2 , and a_3 were initially tabulated by Brosius et al. (1996), and R is the intensity ratio Fe XVI (335.41 Å)/Fe XIV (334.18 Å).

Using the most recent theoretical atomic data from CHIANTI v4.2, the values of a_0, \dots, a_3 were found to change somewhat and resulted in a slightly higher temperature ($\sim 5\%$) than those predicted using Brosius et al. (1996) values. Both previous and updated parameters have been included in Table 4, while the resulting temperatures are presented in Figure 7. In addition, the Fe XIV (334.18 Å) line is density sensitive, and was accounted for by determining the temperature across the active region at the corresponding density.

The temperature maps are displayed in the third row of Figure 5. As expected, the temperature maps show a close spatial correlation with the intensity and density maps. As the region evolves, the temperature remains constant at $\sim 2.25 \times 10^6$ K, with just a slight decrease (by $\sim 4\%$) on the final day of observation. More statistically significant is the variation of electron temperature across the region, which ranges from $\sim 2.10 \times 10^6$ K in the active region boundary to $\sim 2.34 \times 10^6$ K in the core. The values derived for the temperature are heavily dependent on the a_0 coefficient in Equation (8), which accounts for the restricted temperature sensitivity of this method. The temperature is also affected by the brightening on September 13 as can be seen from the corresponding map. Average temperature values are listed in Table 3, and are plotted in Figure 6. Again, these results are in good agreement

with those found in Gallagher et al. (2001) using similar methods.

The C-class flare on September 13 is not only evident in the intensity and density maps of September 13, but was also identified due to the presence of the Fe XVII (347.85 Å) line which has a formation temperature of $\sim 5 \times 10^6$ K. A portion of the spectrum for the flare is shown in Figure 8.

3.4. Electron Pressures

The pressure maps in the bottom row of Figure 5 follow a similar behaviour to that seen in the density. Values for the pressure in the high density core of the active region remain $\sim 10^{16}$ K cm $^{-3}$ from September 10–13 and drop to 4×10^{15} K cm $^{-3}$ on September 14, again not taking into account the effects of the brightening. Similar to the average electron density, the average pressure varied by $\sim 30\%$ between September 10–14 (see the values presented in Table 3 and Figure 6).

Table 4: COMPARISON BETWEEN THE PARAMETERS USED IN EQUATION 8 FOR Fe XVI (335.41 Å)/Fe XIV (334.18 Å).

Study	Log N_e	a_0	a_1	a_2	a_3
Brosius et al. (1996)	9.0	6.237	0.111	0.0087	0.0016
	10.0	6.217	0.108	0.0076	0.0018
Updated values	9.0	6.259	0.111	0.0105	0.0015
	10.0	6.240	0.108	0.0097	0.0015

3.5. Power Balance

The steady-state energetics of a coronal loop can be expressed as,

$$E_H = E_R + E_C + E_F \quad \text{ergs cm}^{-3} \text{ s}^{-1}, \quad (9)$$

where E_R is the total radiative losses of the plasma, E_C is the thermal conductive flux, E_F is the energy lost due to mass flows, and E_H is the energy required to balance these losses (e.g. Antiochos & Sturrock 1982, Bradshaw & Mason 2003).

The radiative loss term, E_R , in Equation (9) can be written in terms of the electron density, N_e , and the radiative loss function, $\Lambda(T_e)$,

$$E_R = -N_e^2 \Lambda(T_e) \quad \text{ergs cm}^{-3} \text{ s}^{-1}. \quad (10)$$

The radiative loss function is usually approximated by analytical expressions of the form $N_e^2 \chi T_e^\alpha$ (Cook et al. 1989). While this method is simple to implement computationally, it does not capture accurately the fine-scale structure of the radiative loss function. A more appropriate technique is used here, which relies on interpolating the values obtained from CHIANTI, using the coronal abundances of Feldman (1992) and the ionization balance calculations of Mazzotta et al. (1998). The choice of coronal abundances was determined by the fact that the Si X ion (used to calculate N_e in Equation (10)) has a formation temperature of 1.26×10^6 K and recent work by Del Zanna & Mason (2003) has shown that active region abundances are coronal.

Assuming classical heat conduction along the magnetic field, the conductive flux can be expressed as,

$$E_C = \frac{d}{ds} \left[-\kappa T_e^{5/2} \frac{dT_e}{ds} \right] \quad \text{ergs cm}^{-3} \text{ s}^{-1}, \quad (11)$$

where the thermal conductivity is $\kappa = 0.92 \times 10^{-6} \text{ ergs s}^{-1} \text{ cm}^{-1} \text{ K}^{-7/2}$ (Spitzer 1962), T_e is the electron temperature. Various approximations for Equation (11) have been used in

previous studies: Aschwanden et al. (1999) used an expression that was heavily dependent upon the temperature gradient, $(dT/ds)^2$, in the study of active region loops, whereas Varady et al. (2000) used an approximation strongly dependent upon temperature, $T^{7/2}$, for post-flare loops. Here, we assume a semi-circular loop geometry and approximate Equation (11) by,

$$E_C = -\kappa T_e^{5/2} \frac{\Delta T_e}{(L/2)^2} \text{ ergs cm}^{-3} \text{ s}^{-1}, \quad (12)$$

where ΔT_e is the difference in the maximum and minimum temperatures given in Section 3.3, and ds is approximated by $L/2$. This approximation is therefore not overly dependent on temperature or the temperature gradient, but acts as a balance between the approximations of Aschwanden et al. (1999) and Varady et al. (2000).

The energy losses due to mass motions of the plasma, E_F , can be expressed as a sum of the kinetic energy and the internal energy of the plasma,

$$E_F = \frac{d}{ds} \left[v \left(\frac{1}{2} \rho v^2 + \frac{5}{2} N_e k_B T_e \right) \right] \text{ ergs cm}^{-3} \text{ s}^{-1}, \quad (13)$$

where v is the flow velocity of the plasma, and ds is again approximated to be the loop half-length ($L/2$). Line-of-sight velocities at the core of the active region were found to be $\sim 10 \text{ km s}^{-1}$ which is consistent with Brynildsen et al. (1999) who detected flow speeds of $\sim 15 \text{ km s}^{-1}$ also using CDS. Using this value, as well as typical parameters from Table 3, the mass loss term is found to be of the order of $\sim 10^{24} \text{ ergs s}^{-1}$ and can therefore be considered negligible compared to radiative losses.

In order to compare the work described in this paper with that of Fisher et al. (1998), Parker (1983), and others, Equation (9) has been integrated over the volume of the active region, and expressed as a power balance equation:

$$P_H = P_R + P_C + P_F \text{ ergs s}^{-1}, \quad (14)$$

where P_R , P_C , and P_F are the power lost due to radiation, conduction, and mass flows,

respectively. P_H is then the heating rate required to balance these losses.

Using the active region properties detailed in Table 3, the radiative, conductive, and mass flow losses were then calculated using Equations (10), (12), and (13). These results, together with the heating rate calculated using Equation (14), are presented in Table 3 and Figure 9. The second panel down of Figure 9 also shows how the conductive flux values depend on how Equation (11) is approximated. In each case, the conductive losses are found to be much less than the radiative losses. Indeed, the average radiative losses ($\langle P_R \rangle \sim 3 \times 10^{25}$ ergs s $^{-1}$) are found to exceed both the conductive losses ($\langle P_C \rangle \sim 2 \times 10^{24}$ ergs s $^{-1}$) and the mass flow losses ($\langle P_F \rangle \sim 4 \times 10^{24}$ ergs s $^{-1}$) by approximately an order of magnitude. Aschwanden et al. (2000) also found an order of magnitude difference between the radiative and conductive losses in coronal loops despite the actual values being somewhat lower than those found here due to the insensitivity of EIT filter ratio techniques. As can be seen from Figure 9, the heating rate falls by close to a factor of 5 between September 10 and 14 and is not significantly affected by whichever conductive loss equation is used.

Figure 10 shows the heating rate as a function of the total unsigned magnetic flux. A least-squares fit to the non-flaring data (i.e., neglecting the high density plasma of September 13) yielded a power-law of the form $P_H \sim \Phi_{tot}^\gamma$, where $\gamma = 0.81 \pm 0.32$. Power-laws with slopes of 1 and 2 from Equations (1) and (3) are also shown in Figure 10 for comparison with our own fit.

4. DISCUSSION AND CONCLUSIONS

A detailed study of an evolving active region has been described using measurements from several instruments onboard *SOHO*. The region was observed to decay in size and complexity as it passed from close to the central meridian on September 10, to close to the

west limb on September 14. In the photosphere, the total sunspot area fell by close to a factor of 2, while the total magnetic flux fell by approximately a factor of 6. In the corona, the average electron density, temperature, and pressure all showed similar decreases in value, which is to be expected considering that the plasma in the corona traces out field lines which are ultimately rooted in the photosphere. These results are of consequence to efforts in understanding active region evolution and the relationship between the photosphere and corona of solar active regions (e.g., Abbett & Fisher 2003; Ryutova & Shine 2004).

In addition to studying active region evolution, CDS and MDI were used to investigate the power-law relationships predicted by theoretical models of the corona. Mandrini et al. (2000) also investigated theoretical scaling laws using magnetic field extrapolations from both vector and longitudinal magnetograms, in light of the work of Klimchuk & Porter (1995). They concluded that models involving the gradual stressing of the coronal magnetic field are in better agreement with the observational constraints than are wave heating models. This same general conclusion was also reached by Démoulin et al. (2003) using photospheric and coronal measurements from MDI and *Yohkoh*. Unfortunately these studies relied on broad-band filter ratios to determine electron temperatures and densities; the difficulty associated with making such measurements is clear from the contradictory results of Priest et al. (1998), Aschwanden (2001), and Reale (2002), who also investigated coronal loop heating models using filter ratios from *Yohkoh*/SXT.

The analysis presented in this paper, on the other hand, is based on well understood line ratio techniques, which offer a less ambiguous determination of plasma properties. With this in mind, the power-law relationship between the total heating rate and the total unsigned magnetic flux (Φ_{tot}) was determined, finding a relationship of the form $P_H \sim \Phi_{tot}^{0.81 \pm 0.32}$. Fisher et al. (1998) compared the X-ray luminosity (which was assumed to be some fraction of the total heating power) to active region vector magnetograms, finding a similar power-law

relation of $L_X \sim \Phi_{tot}^{1.19}$. The result of Fisher et al. (1998) suggests a “universal” relationship between magnetic flux and the amount of coronal heating, regardless of the age or complexity of the active region. A similar relationship of $L_X \sim \Phi_{tot}^{0.9}$ was found using statistical samples of late-type stars (Schrijver 1987). The results of this paper therefore lend further observational evidence that active regions are heated by magnetically-associated waves, rather than multiple nanoflare-type events.

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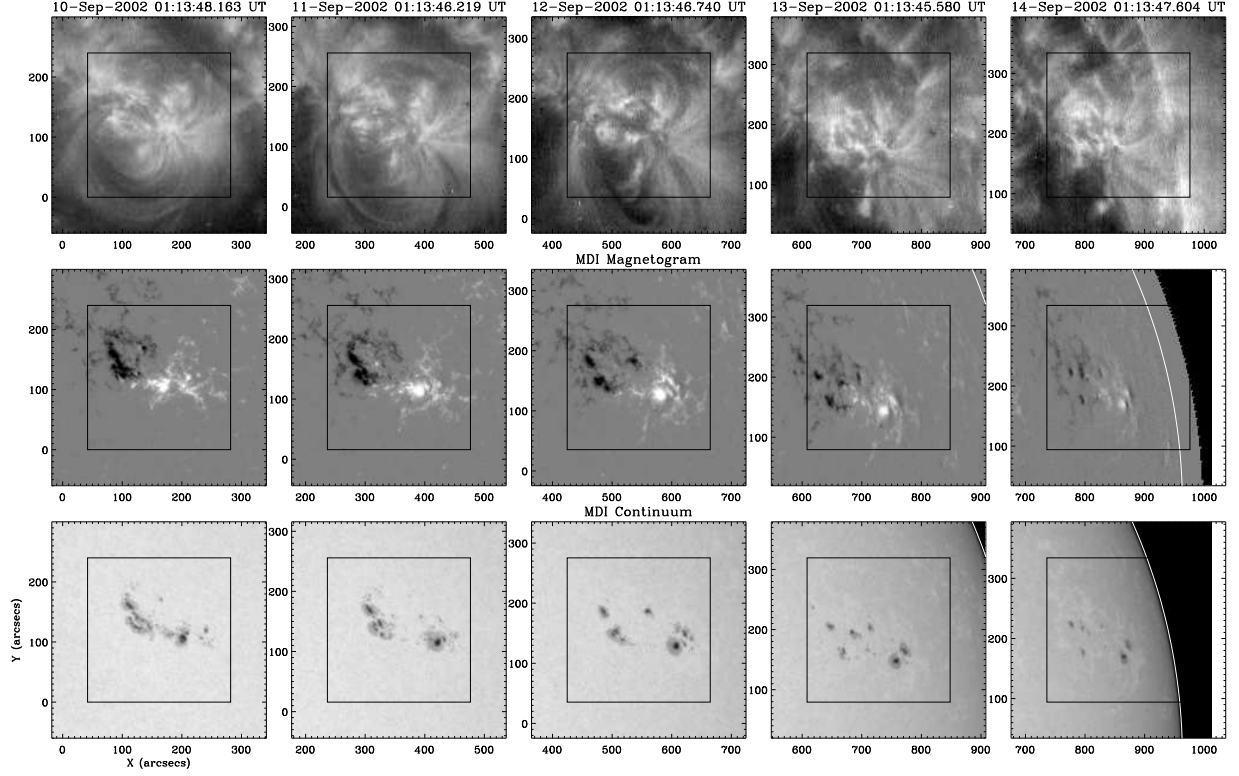


Fig. 1.— The evolution of NOAA 10103 during September 10–14. *Top row:* EIT (195 Å) images. *Middle row:* MDI magnetograms. *Bottom row:* MDI continuum images. The inner boxes in each image denote the field-of-view of the CDS raster.

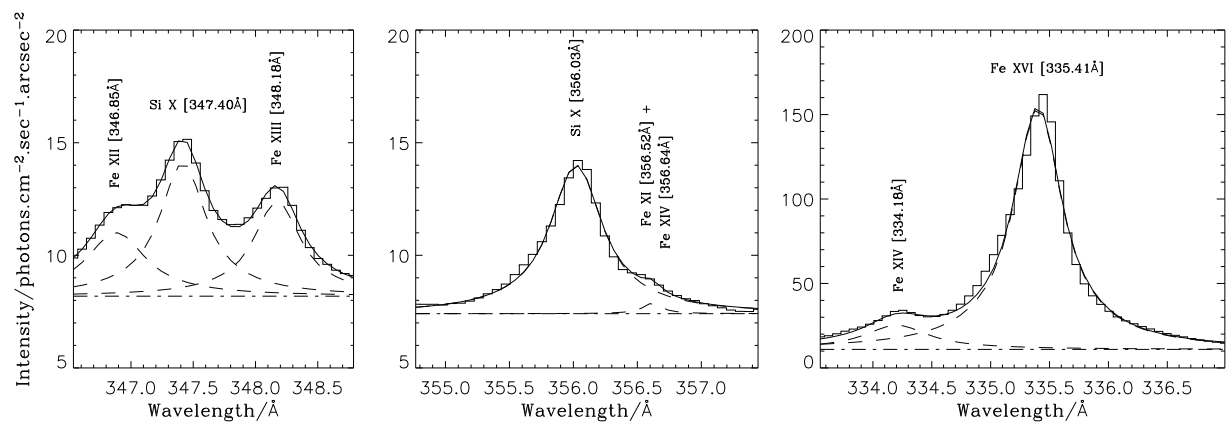


Fig. 2.— Portions of the CDS spectra from NOAA 10103 on September 10, together with backgrounds (horizontal dot-dash lines) and broadened Gaussian fits.

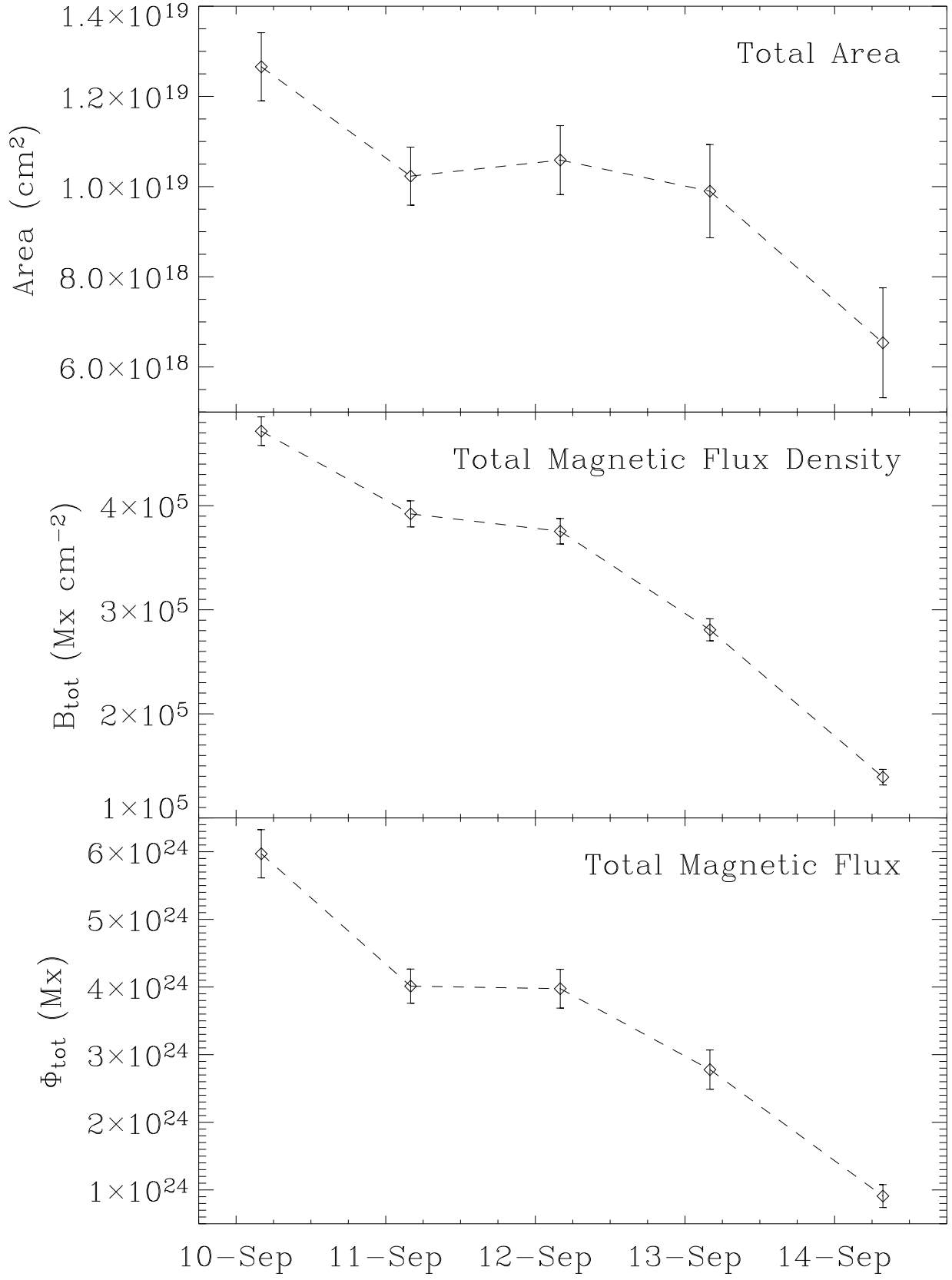


Fig. 3.— The cosine-corrected active region area, total magnetic flux density, and total magnetic flux for September 10–14.

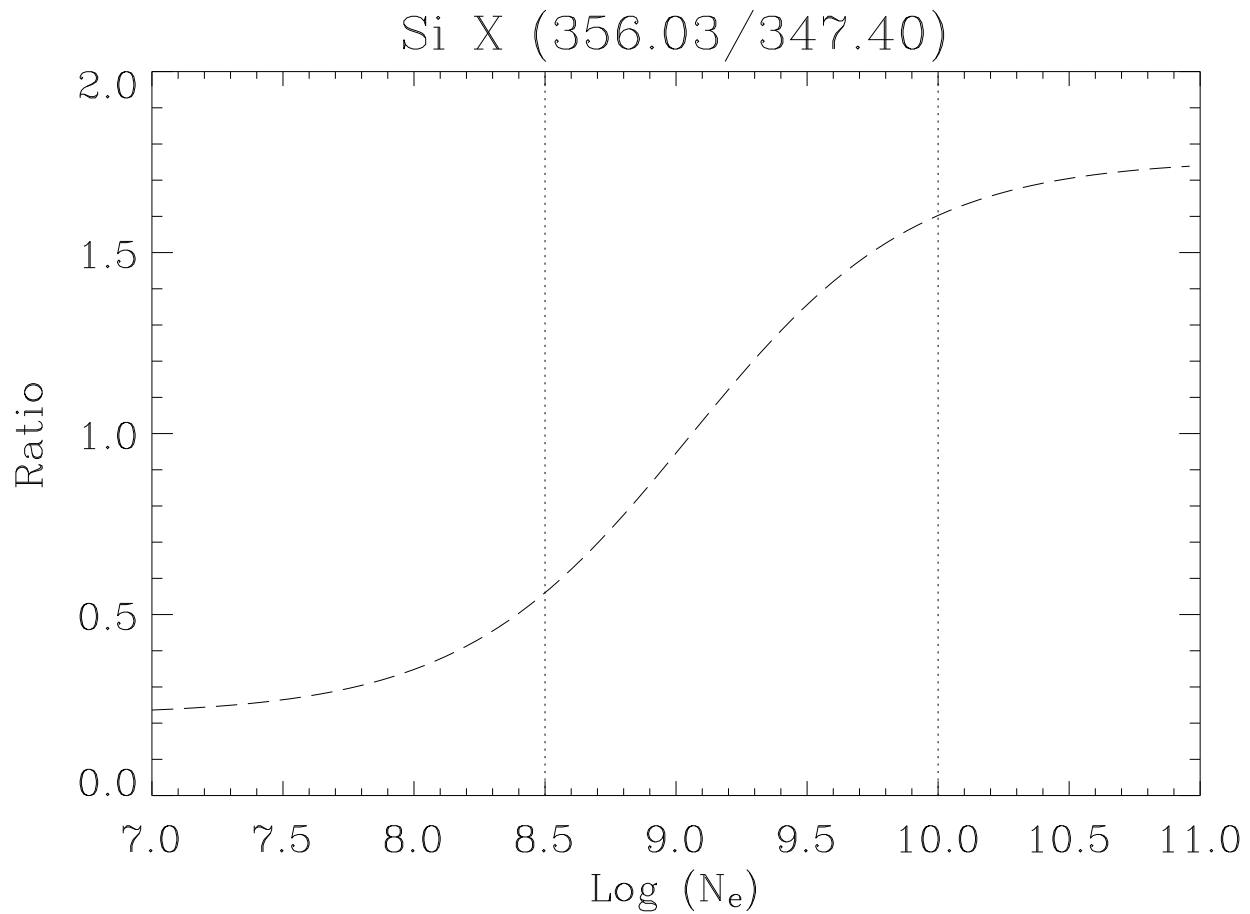


Fig. 4.— Theoretical line ratio as a function of density from CHIANTI for Si X (356.03 Å/347.41 Å). The vertical dotted lines denote the typical range of densities expected to be found in an active region.

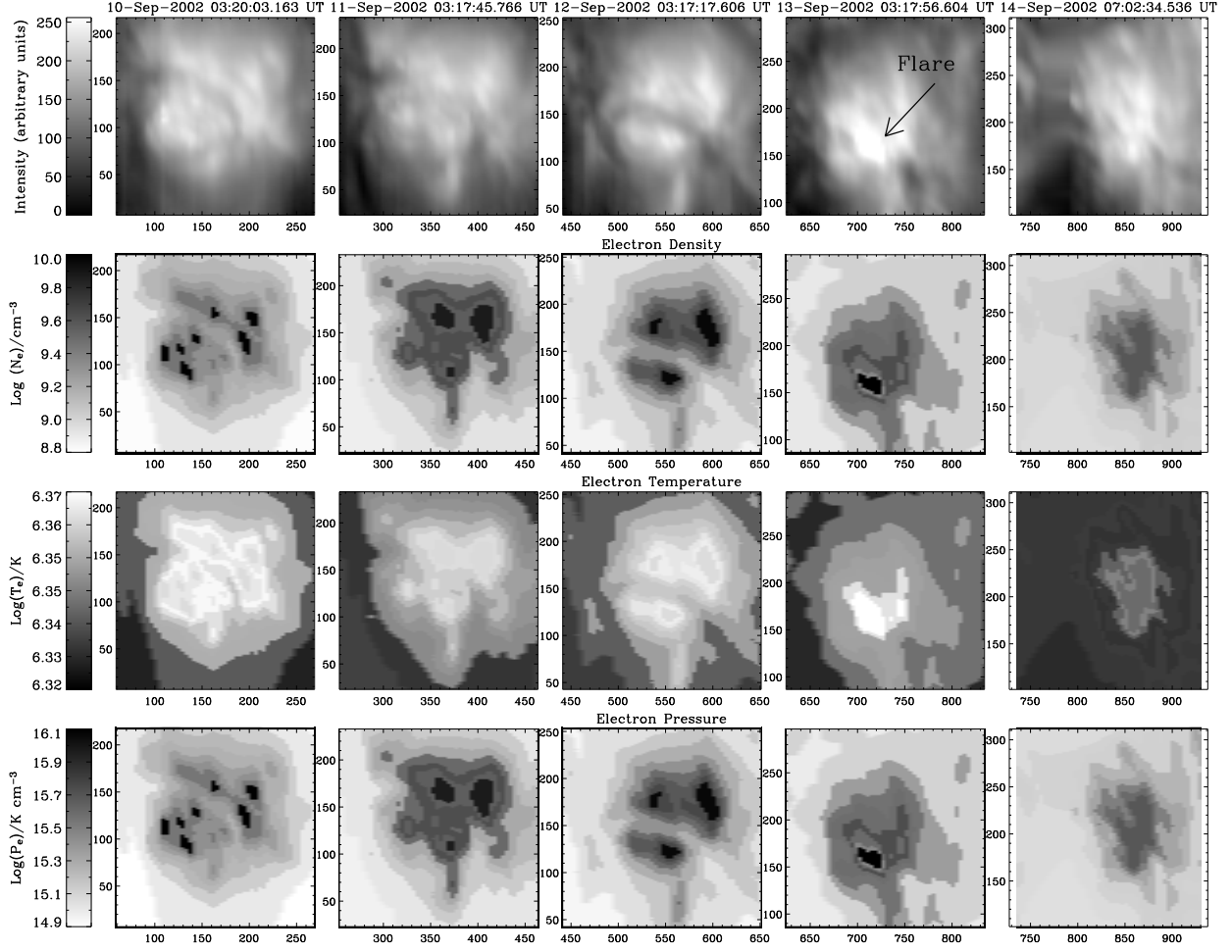


Fig. 5.— *Top row:* CDS images of NOAA 10103 over the five days of observations. *Second row:* Electron density maps derived from the Si X (356.03 Å/347.41 Å) ratio. *Third row:* Electron temperature maps found from the Fe XVI (335.41 Å)/Fe XIV (334.19 Å) ratio. *Bottom row:* Electron pressure maps from the product of the density maps with the formation temperature of the Si X lines. The panels for September 13 show the compact brightening.

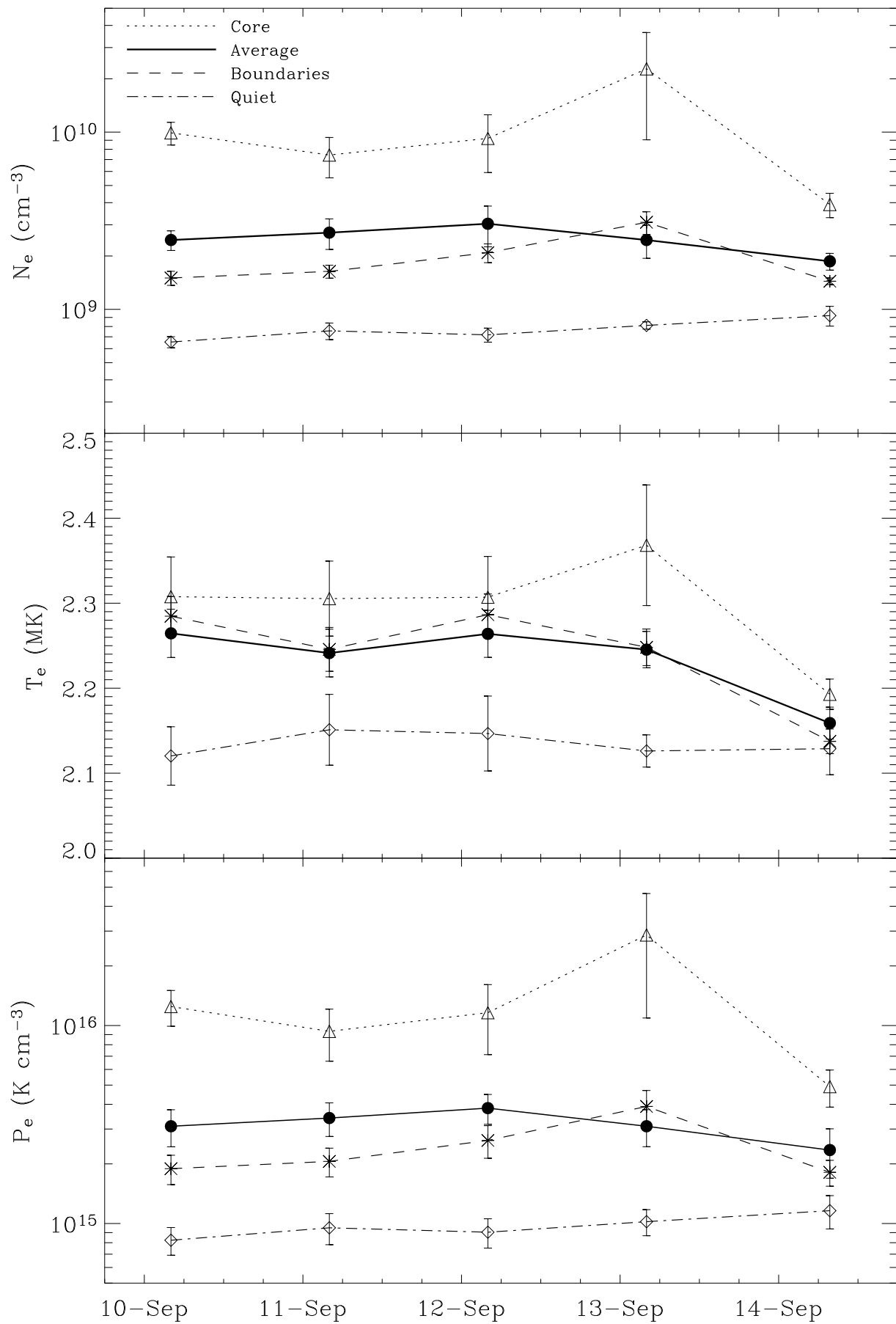


Fig. 6.— Variations in electron density (N_e), temperature (T_e), and pressure (P_e) for September 10–14.

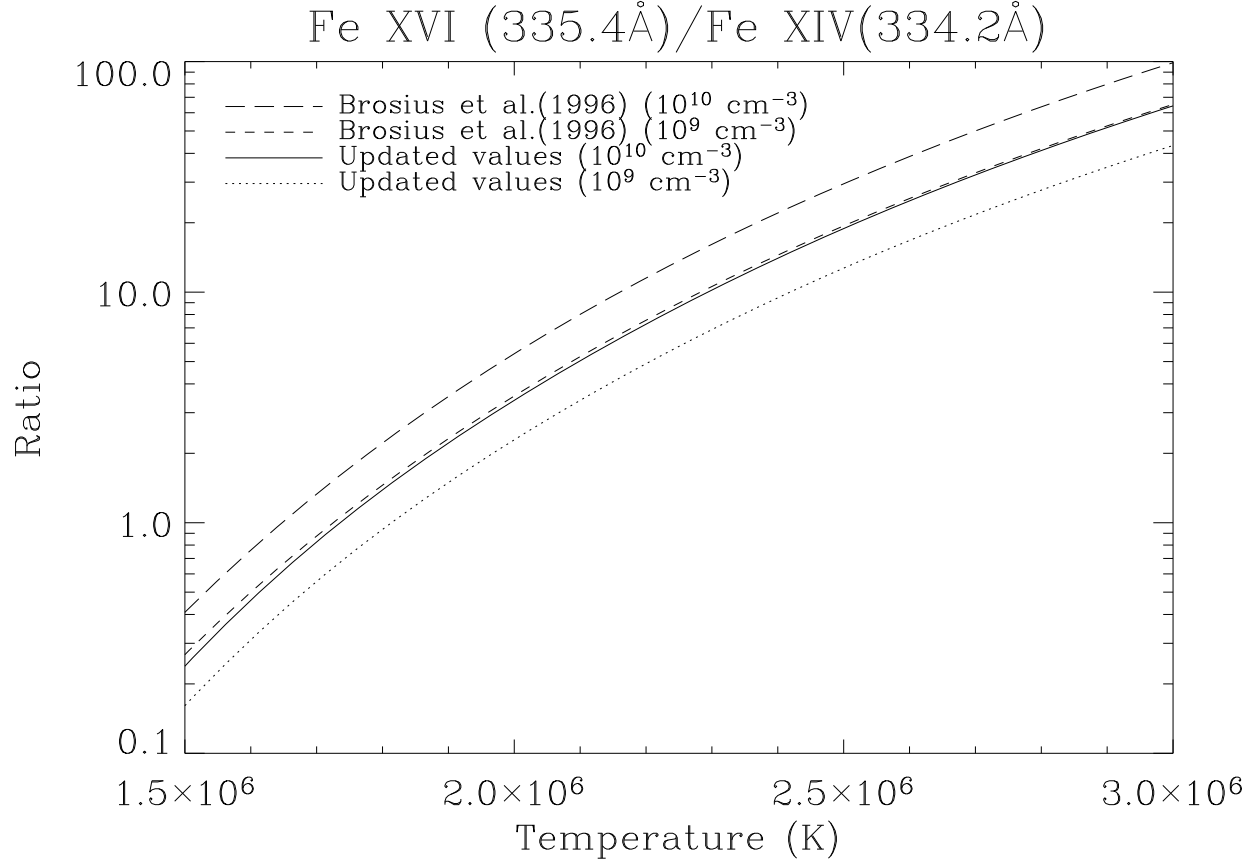


Fig. 7.— The temperature sensitive curves for Fe XVI (335.41 Å)/Fe XIV (334.18 Å) for the Brosius et al. (1996) method and the updated CHIANTI v4.2 data, both plotted for plasma densities of 10^9 and 10^{10} cm^{-3} .

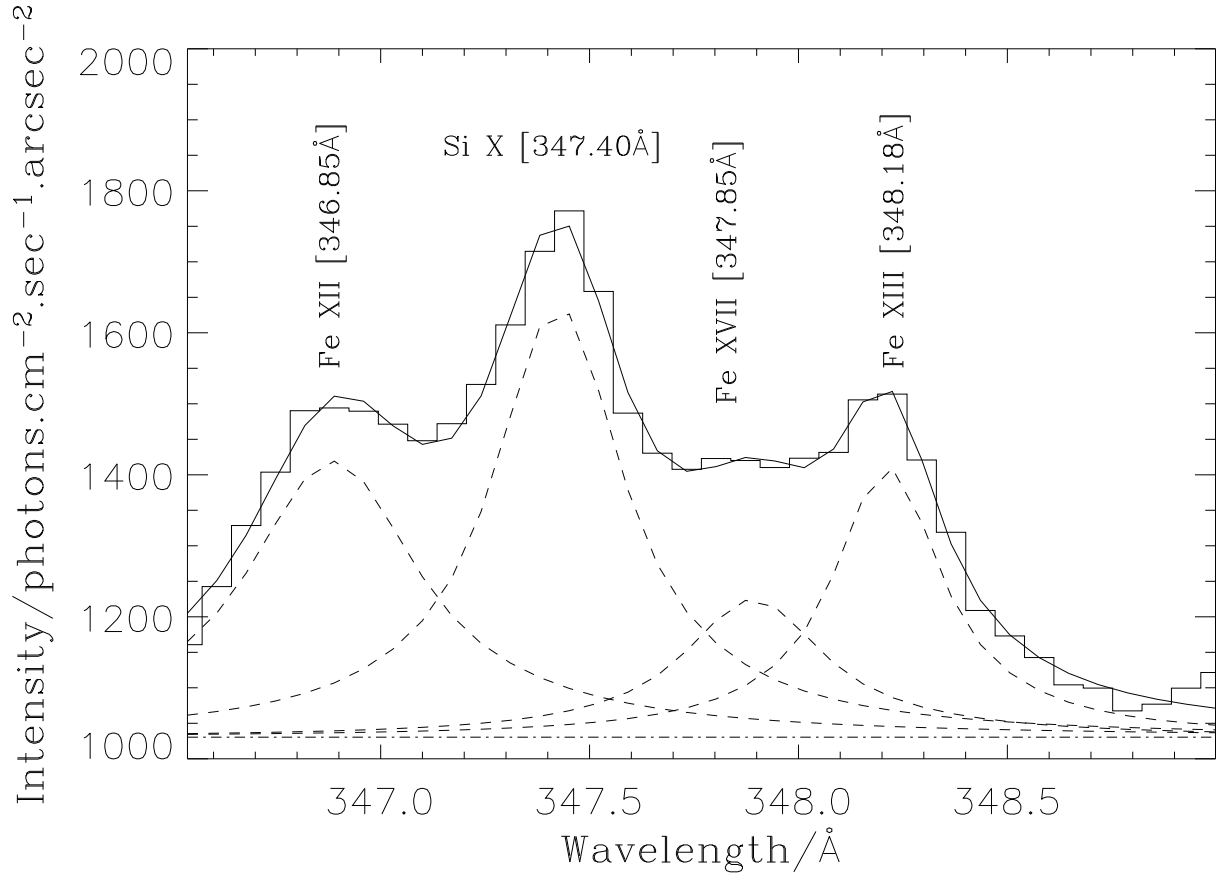


Fig. 8.— Portion of the spectrum from the compact brightening on September 13 showing the Fe XVII (347.85 Å) emission line and neighbouring lines. Fe XVII has a formation temperature of $\sim 5 \times 10^6$ K.

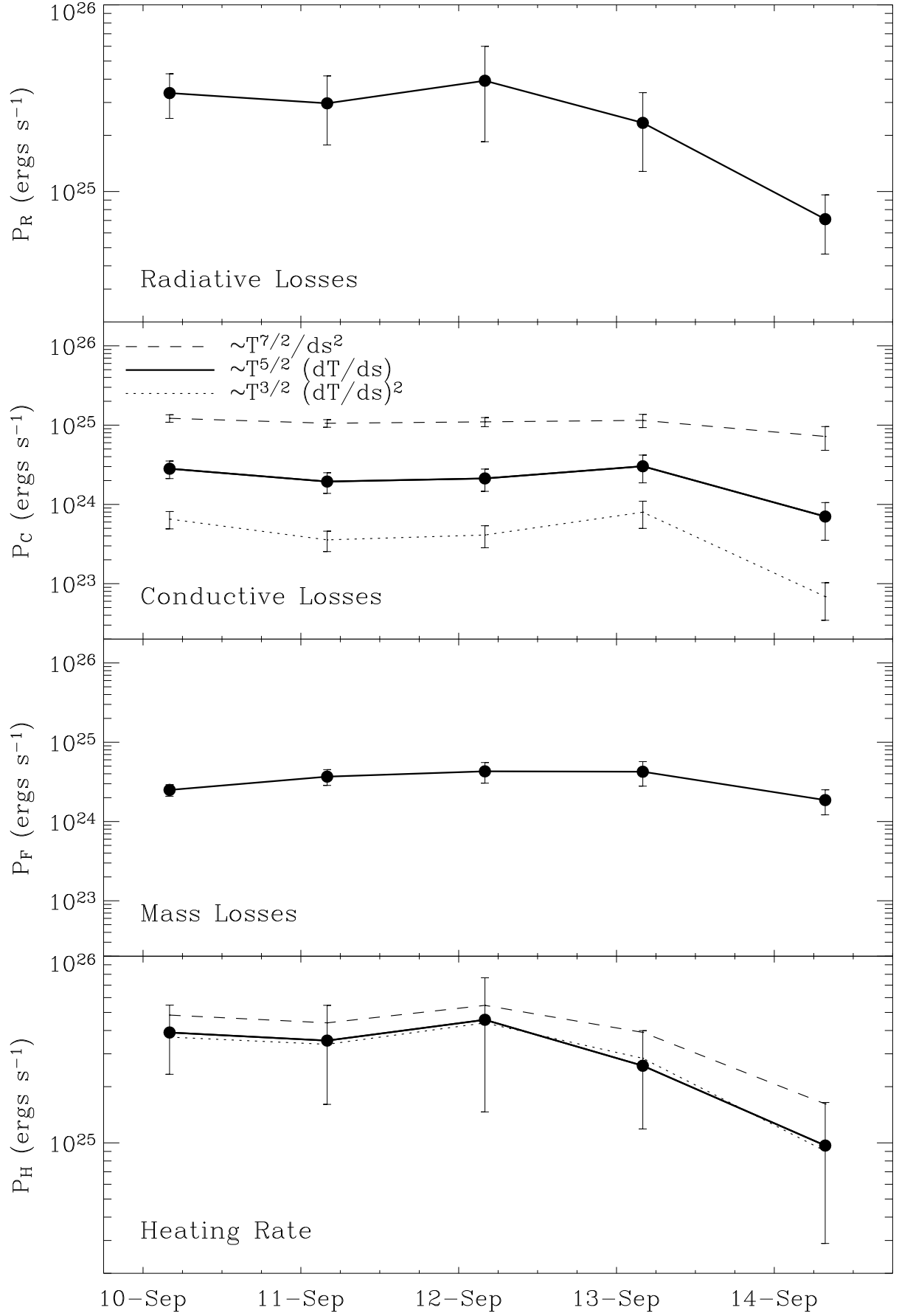


Fig. 9.— The radiative, conductive, and mass losses, and the heating rate for NOAA 10103 for September 10–14.

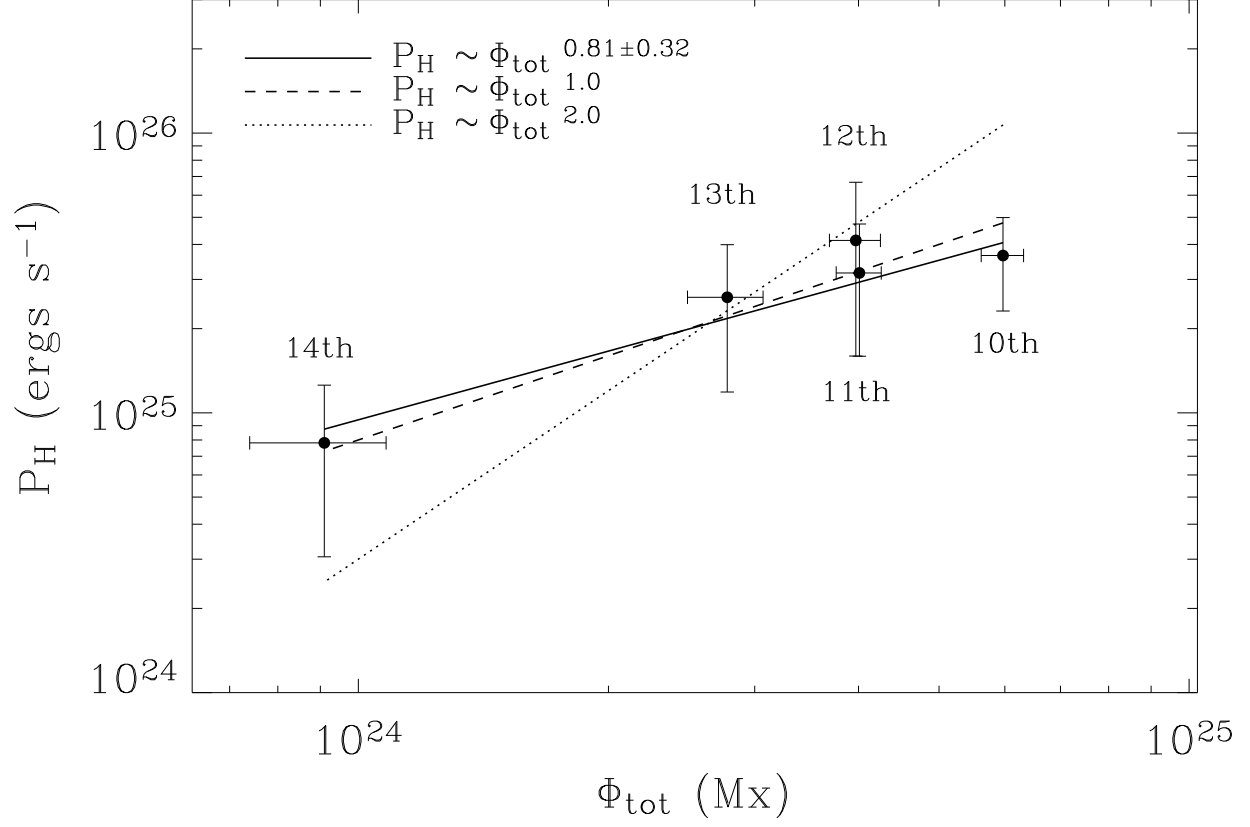


Fig. 10.— Plot of P_H versus Φ_{tot} for five days of MDI and CDS observations. Also shown is a least-squares fit to the non-flaring data together with power-laws with slopes of 1 and 2 from the theoretical predictions.